

# A new study of the chemical structure of the Horsehead nebula: the influence of grain-surface chemistry

**Romane Le Gal**

Post-doctoral fellow (CfA-Harvard)  
in Karin Oberg's group

**Collaborators for this work:**

E. Herbst (UVa), G. Dufour (NASA Goddard),  
P. Gratier (LAB), M. Ruaud (NASA Ames),  
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**Astronomy  
&  
Astrophysics**

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<sup>4</sup> NASA Ames Research Center, Moffett Field, CA 94035, USA



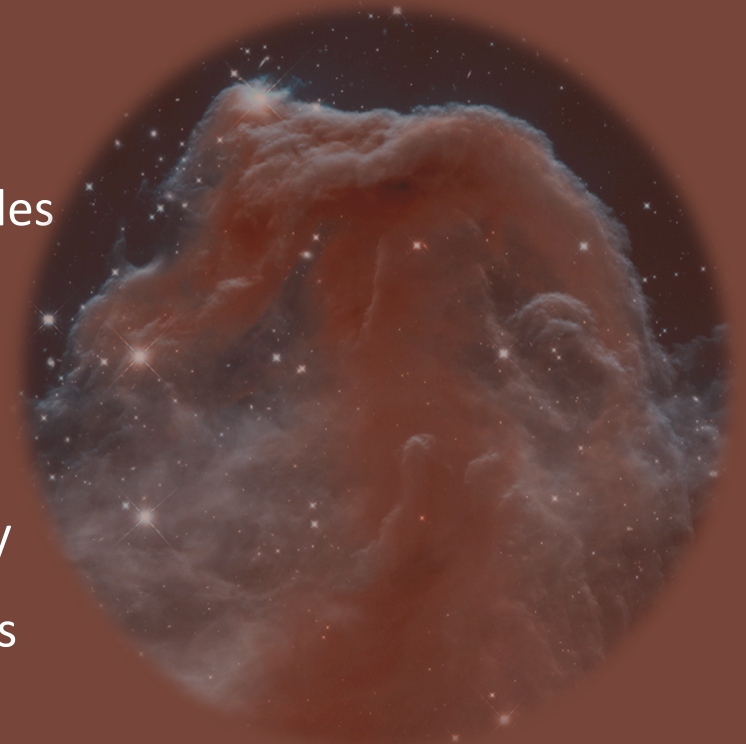
# Outline

- **Context**

- Detection of a wide variety of molecules in the Horsehead nebula

- **Our study**

- Model developed: gas-grain chemistry
- Results: comparison with observations

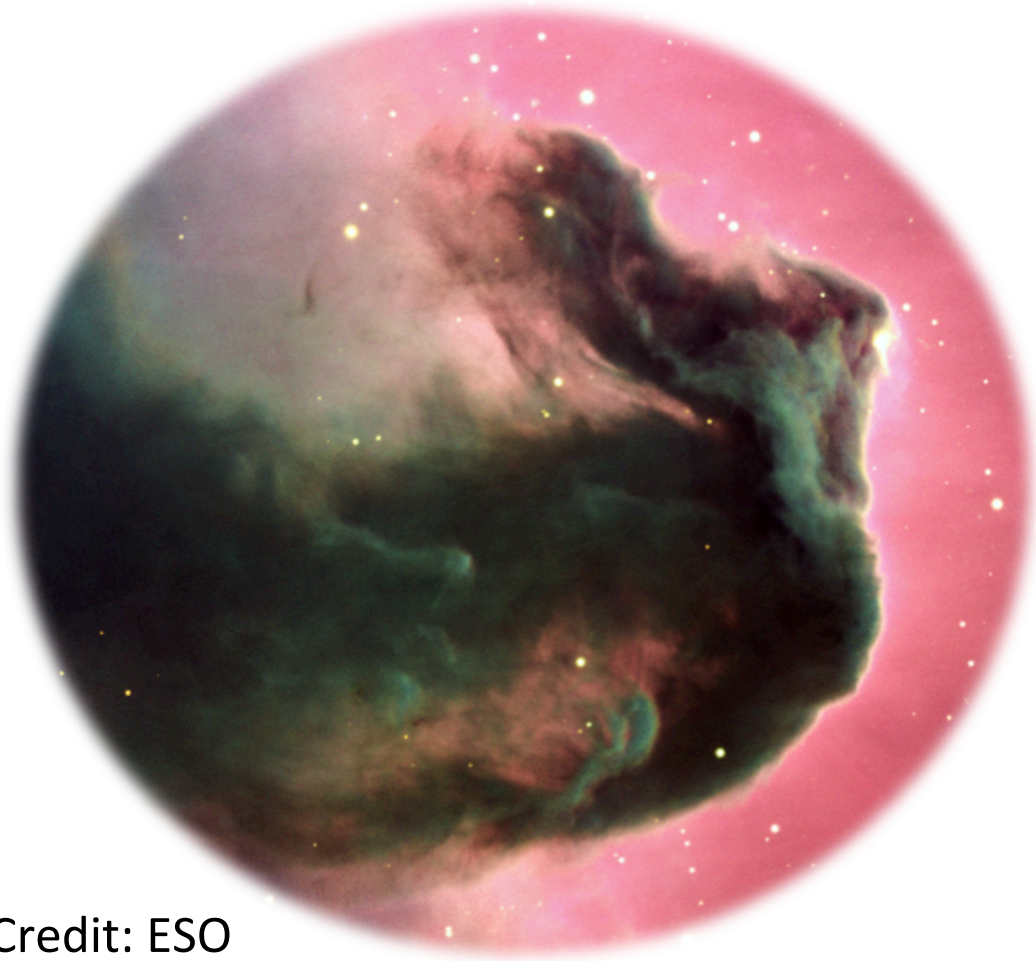


# The Horsehead nebula





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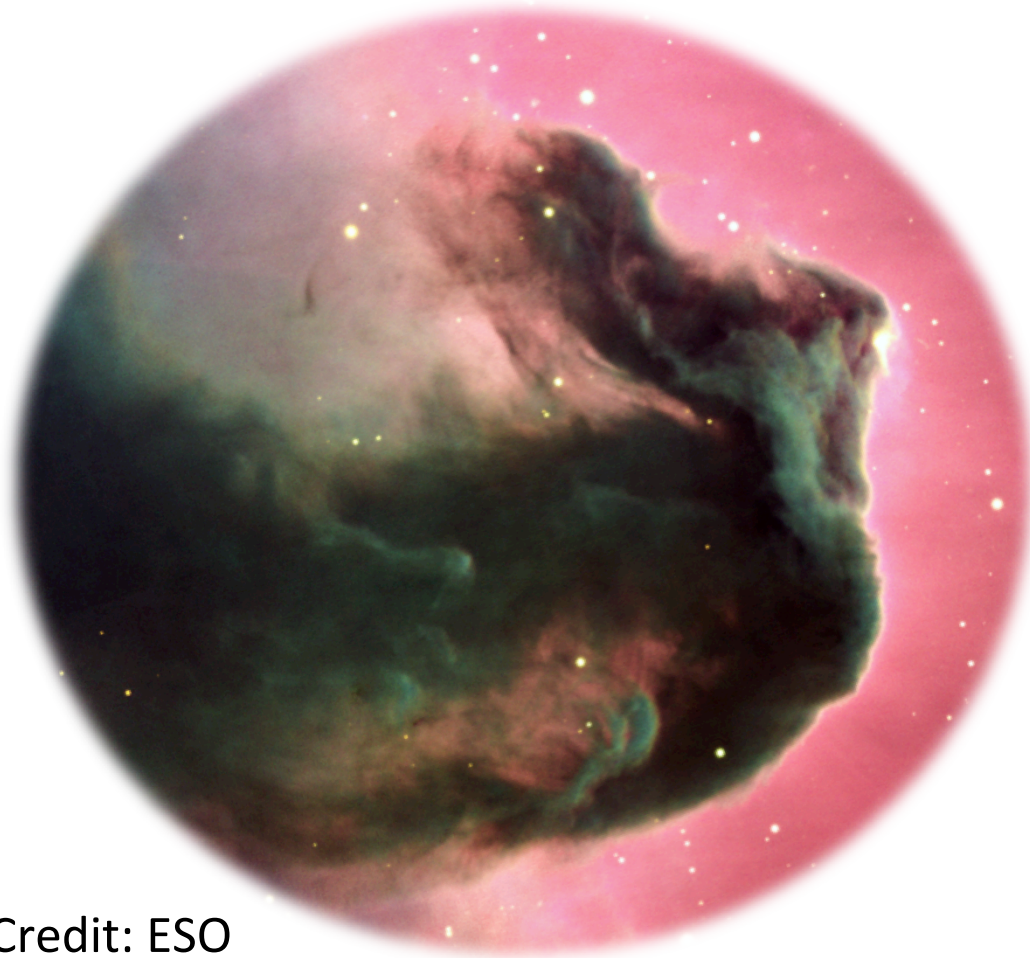


Credit: ESO

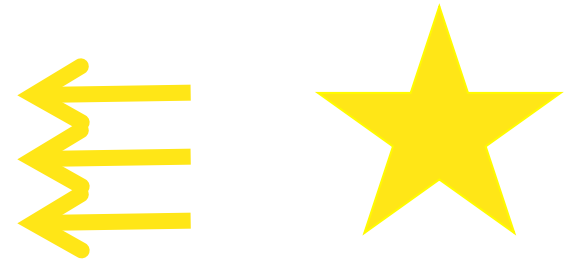
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Nearly edge-on (Abergel et al. 2003)

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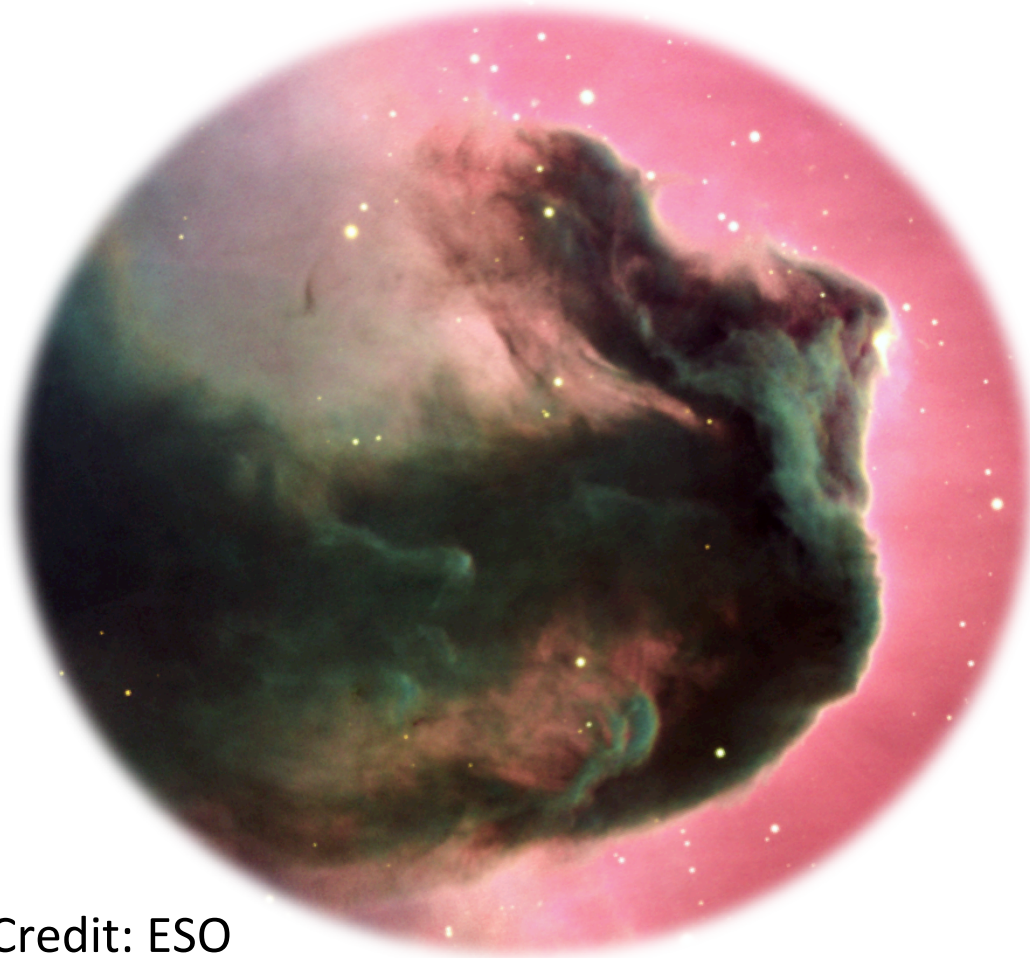


$\sigma$  Orionis  
O9.5 star

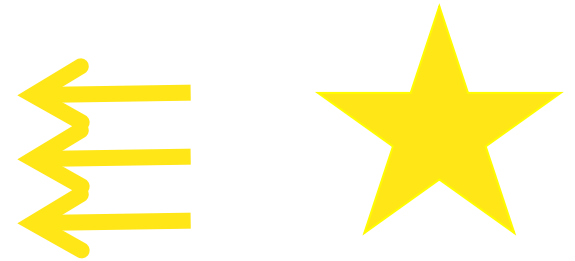
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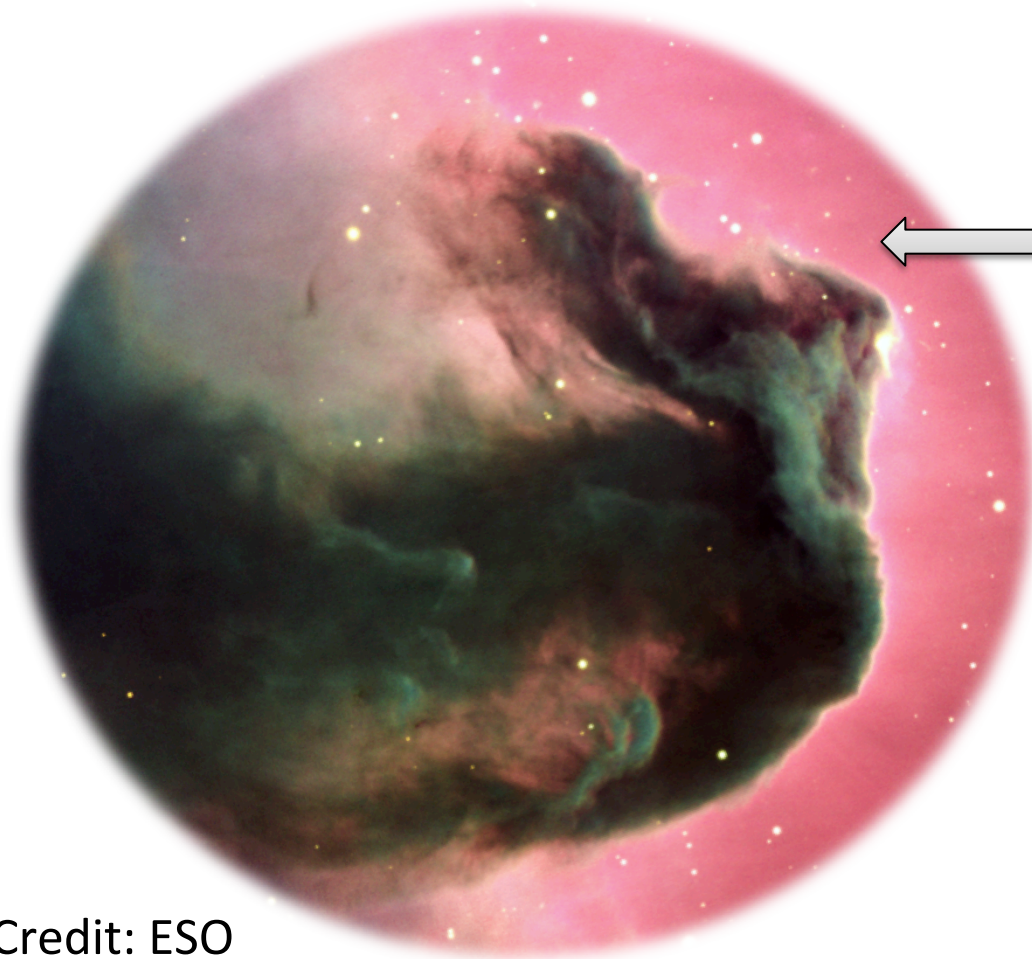
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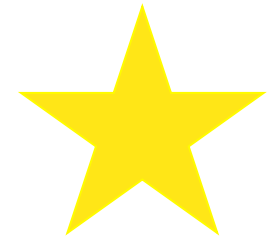
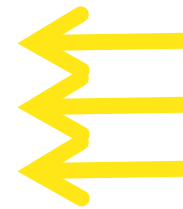
with  $\chi = 60$  x ISRF (Mathis et al. 1983; Habart et al. 2005)



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←  $\approx 3.5$  pc →



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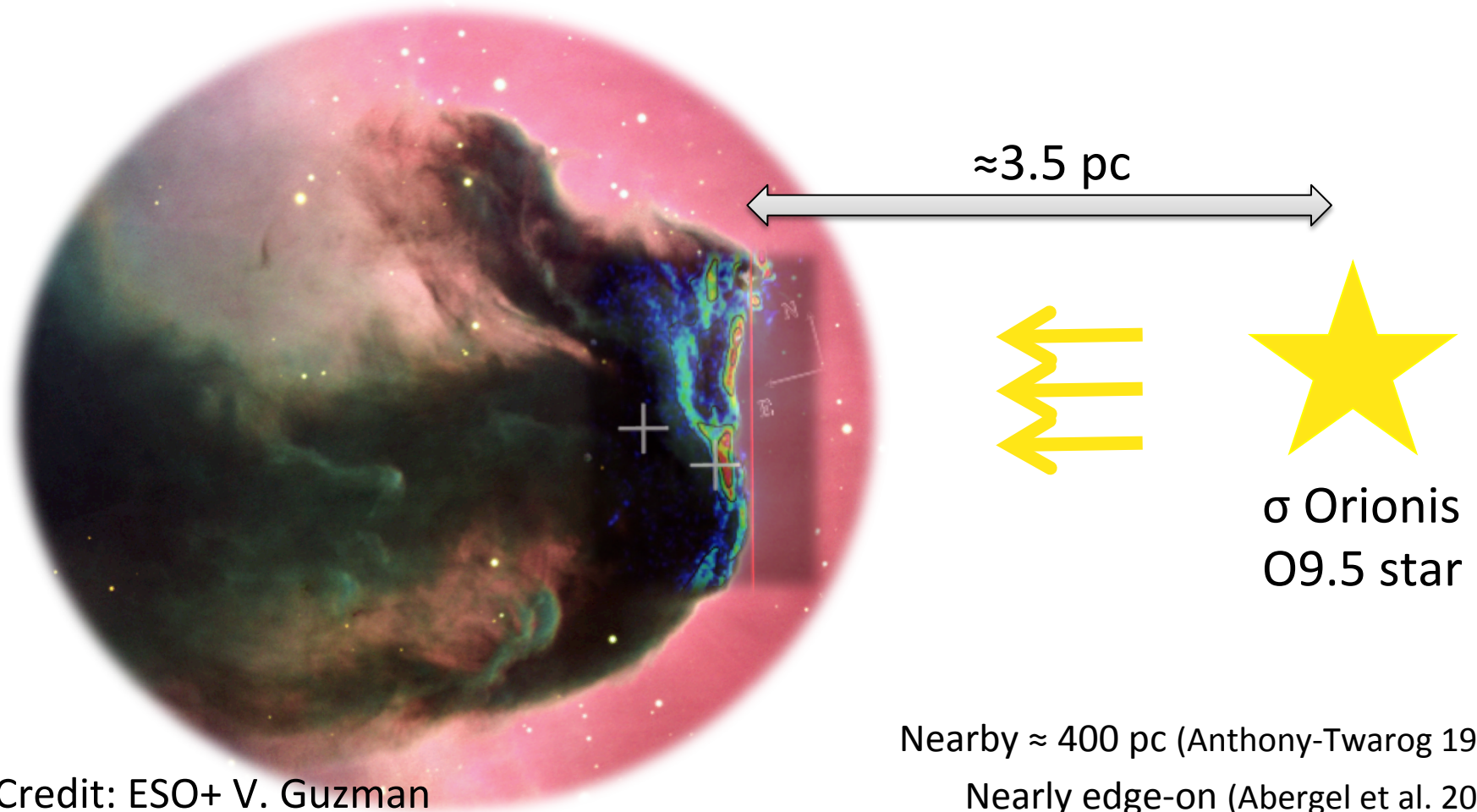
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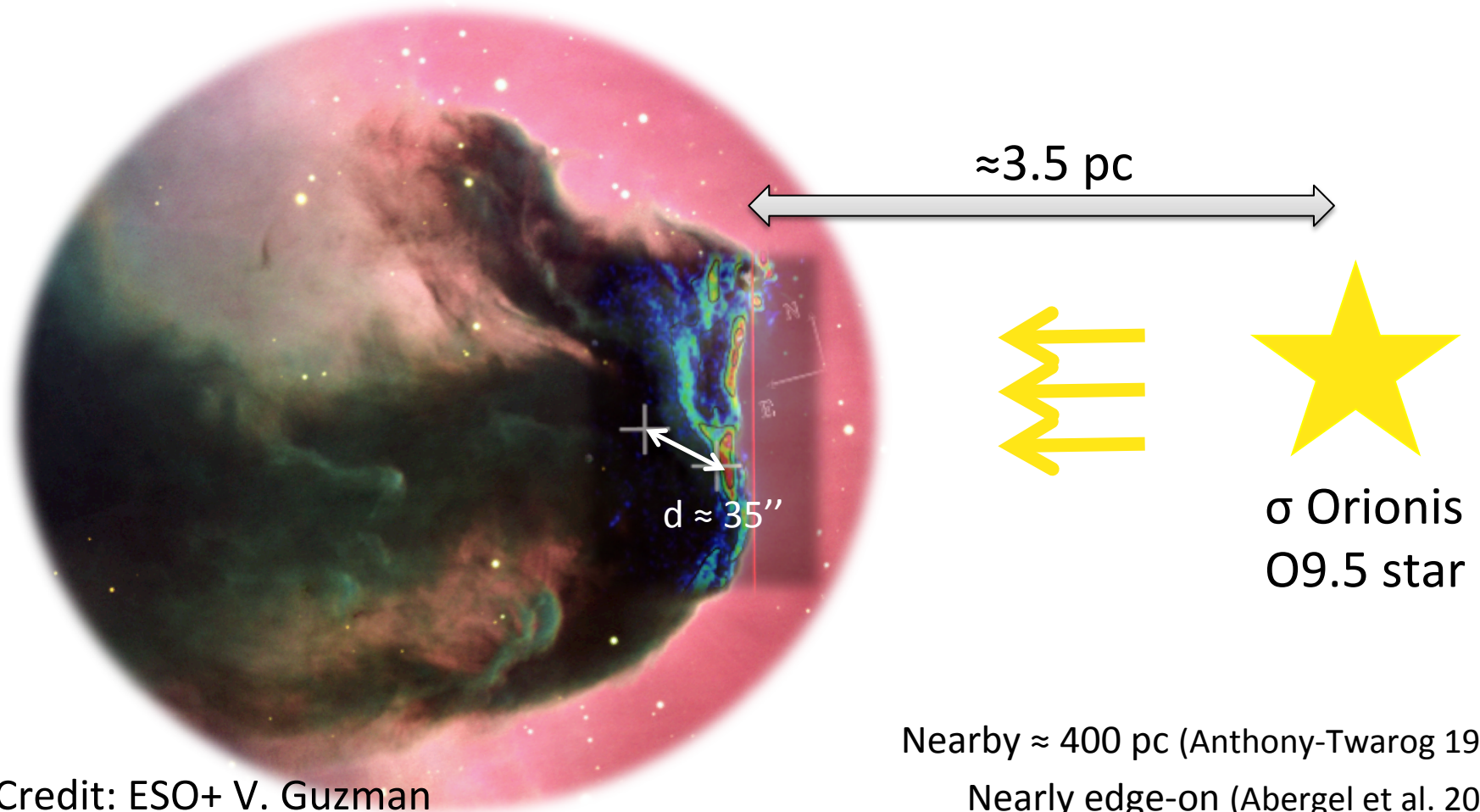
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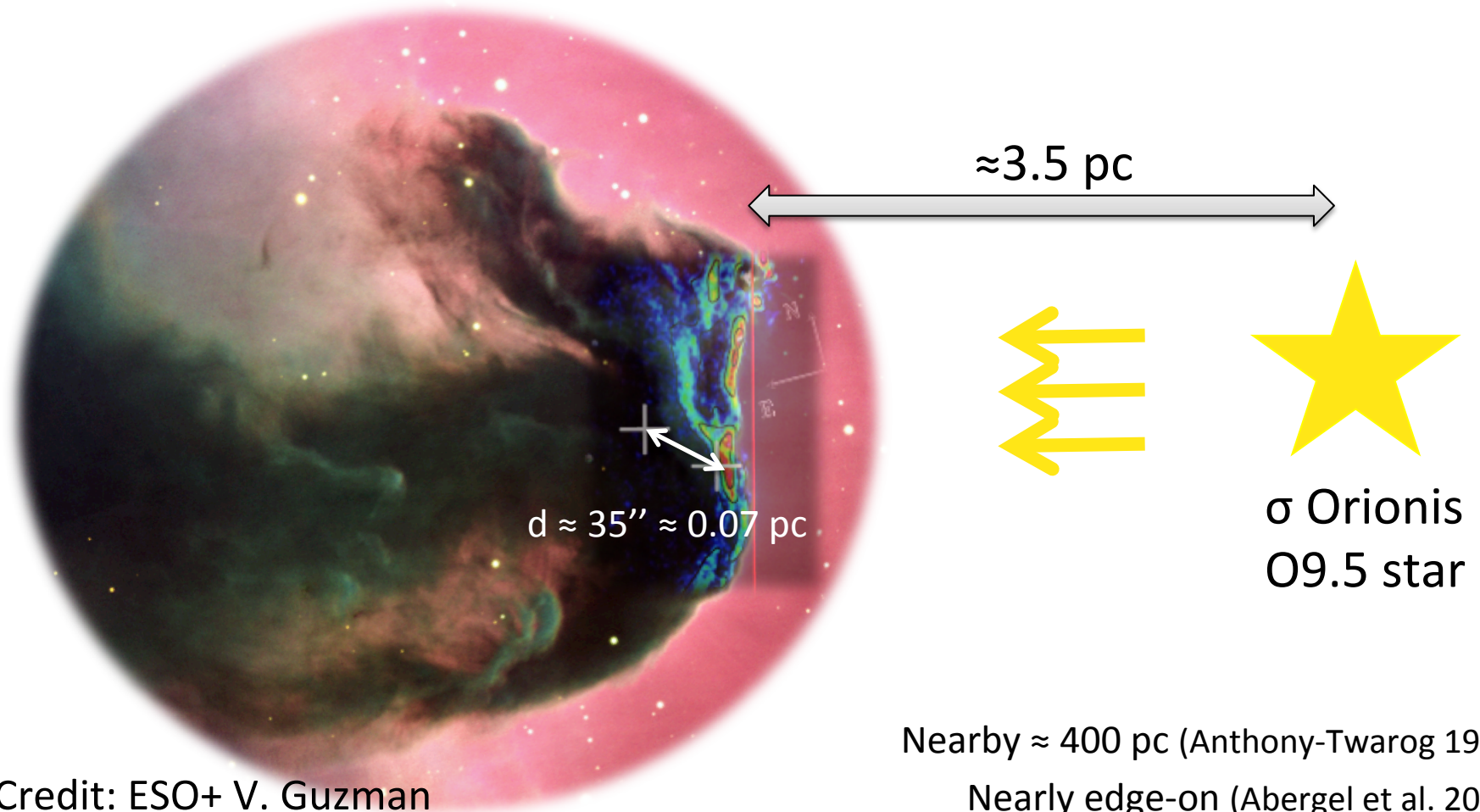
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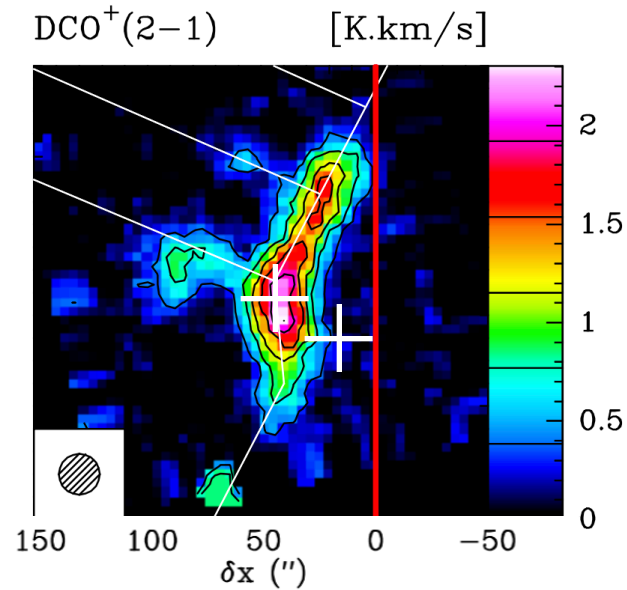
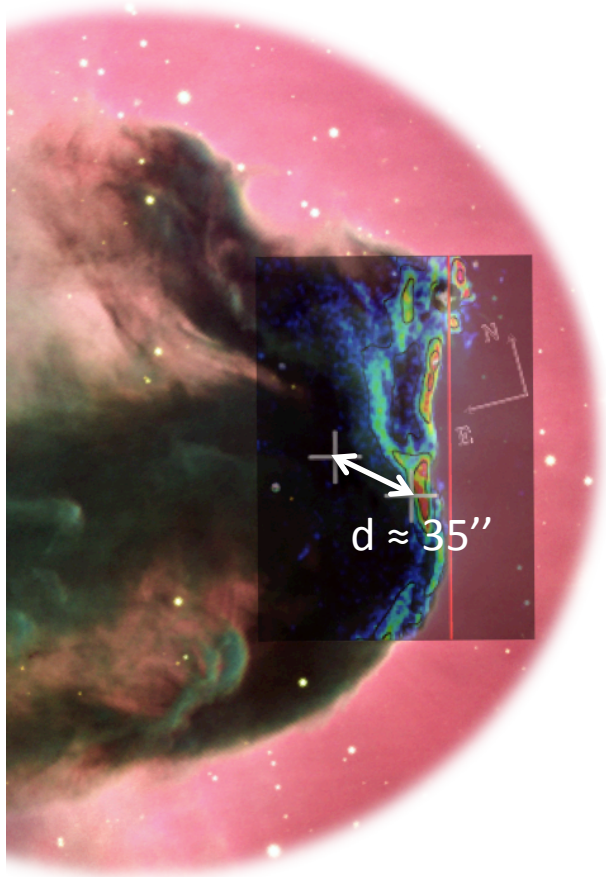
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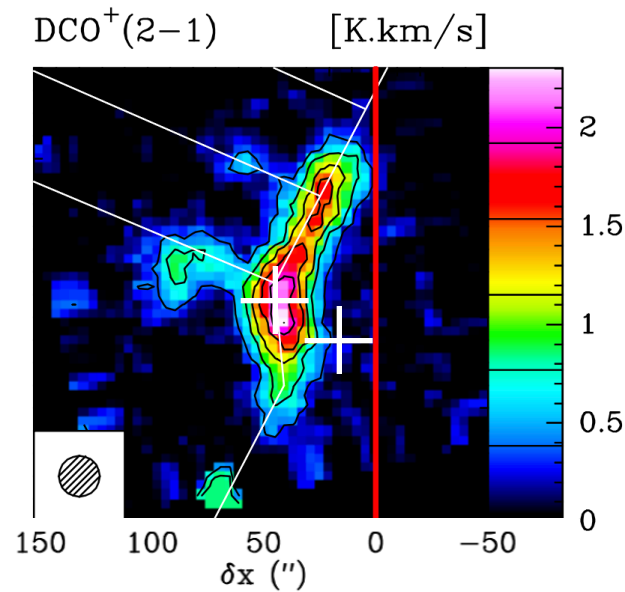
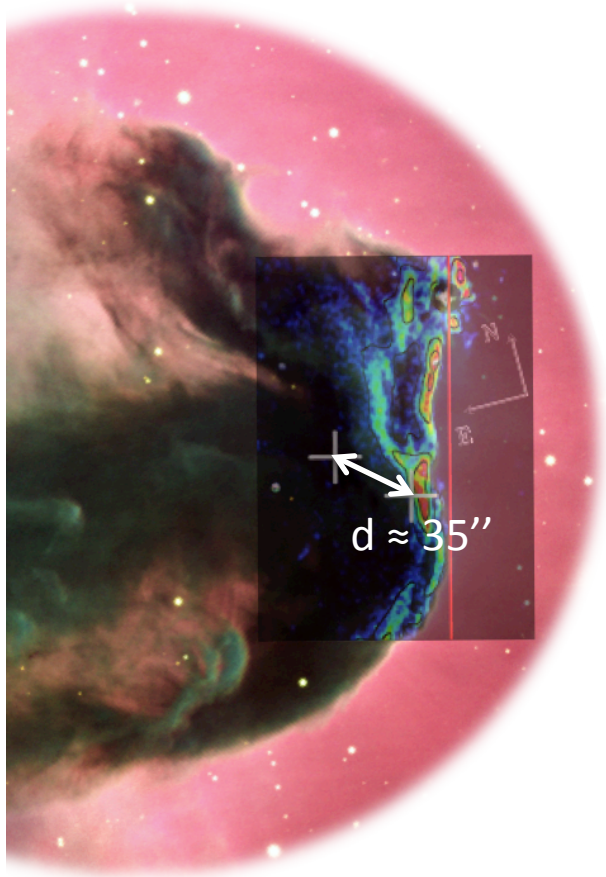
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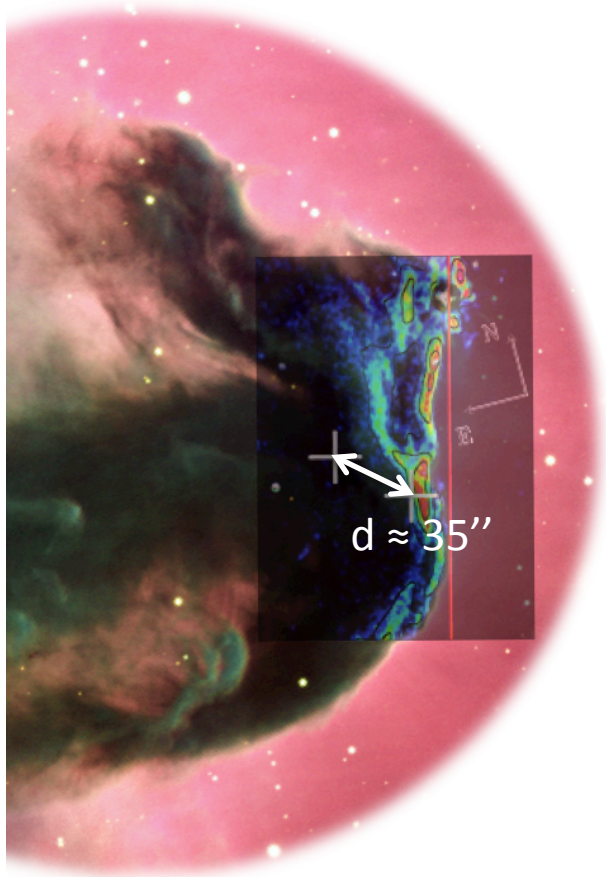


A shielded, dense  
core (Pety+ 2007)  
 $T \approx 20$  K

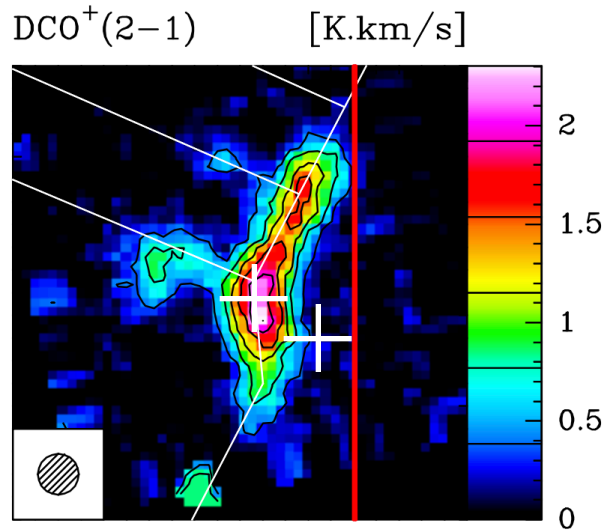
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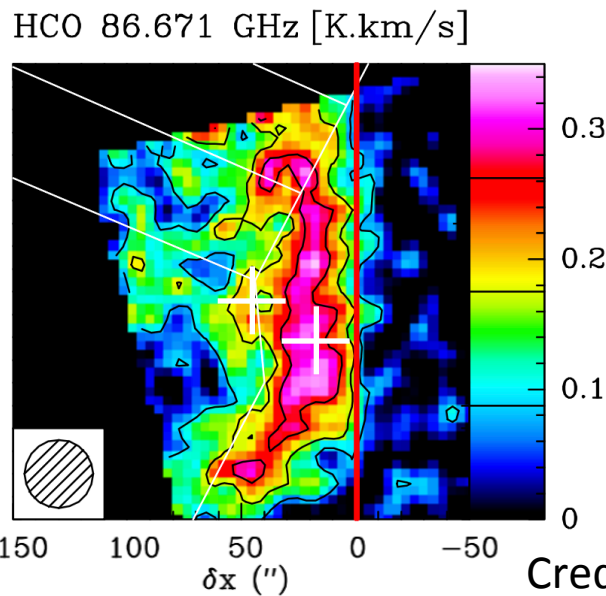
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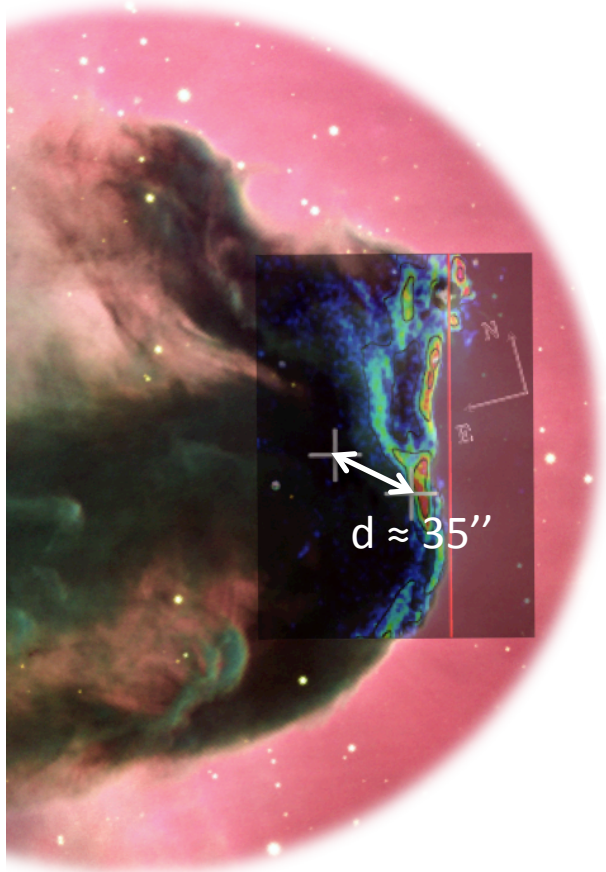
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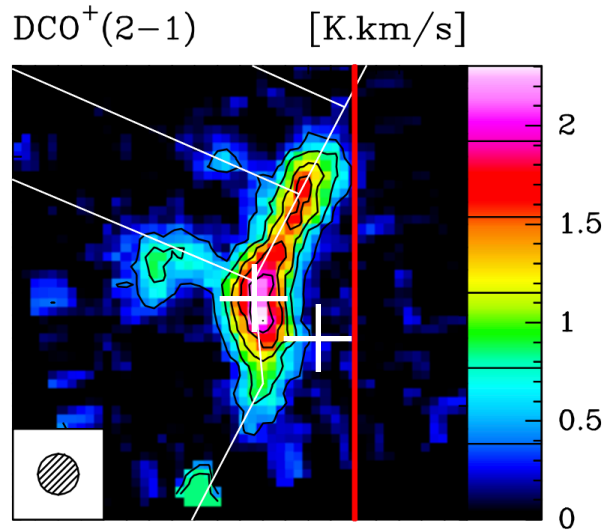
Credit: J. Pety HDR



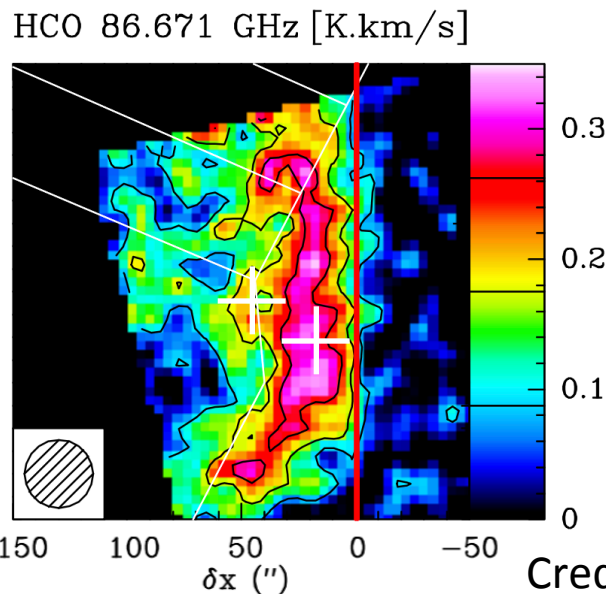
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Credit: ESO+ V. Guzman



A shielded, dense core (Pety+ 2007)  
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A far-UV illuminated PDR (Gerin+ 2009)  
 $T \approx 60$  K

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Organic molecules and precursors	Nitriles	Small hydrocarbons	F-bearing molecules
HCO	CH <sub>3</sub> CN	CCH	CF <sup>+</sup>
H <sub>2</sub> CO	HC <sub>3</sub> N	l-C <sub>3</sub> H	
CH <sub>3</sub> OH	C <sub>3</sub> N	c-C <sub>3</sub> H	
HCOOH		l-C <sub>3</sub> H <sub>2</sub>	
CH <sub>2</sub> CO		c-C <sub>3</sub> H <sub>2</sub>	
CH <sub>3</sub> CHO		l-C <sub>3</sub> H <sup>+</sup>	
CH <sub>3</sub> CCH			



# A new astrochemical model

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- Detection of a wide variety of molecules in the Horsehead nebula



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chemical evolution during  $1e6$  yr  
of a starless dense cloud with  
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$T = 10 \text{ K},$   
 $n_{\text{H}} = 2 \times 10^4 \text{ cm}^{-3},$   
 $A_{\text{V}} = 30 \text{ mag},$   
 $\zeta = 5 \times 10^{-17} \text{ s}^{-1}$



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Chemical  
network of  
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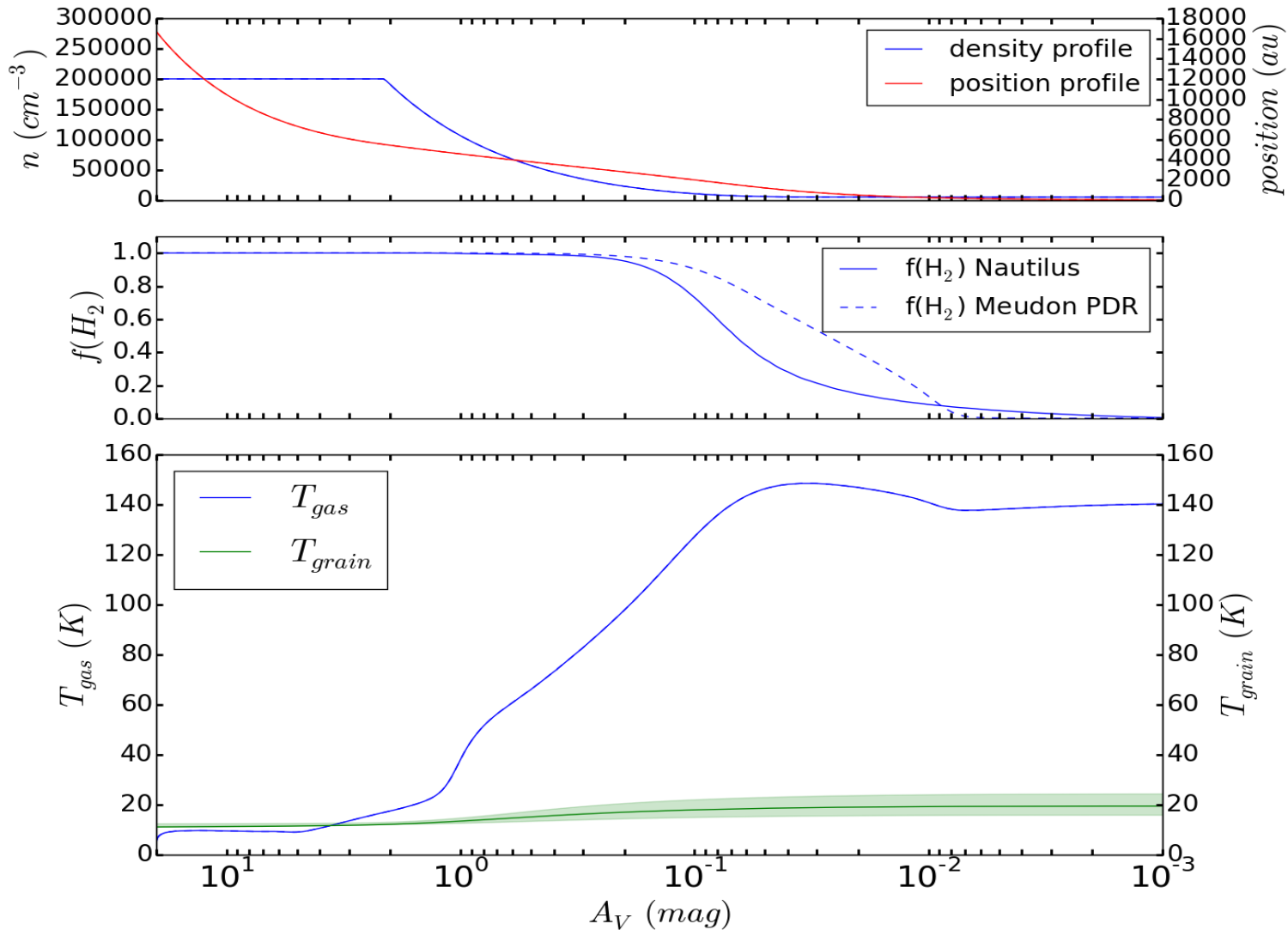
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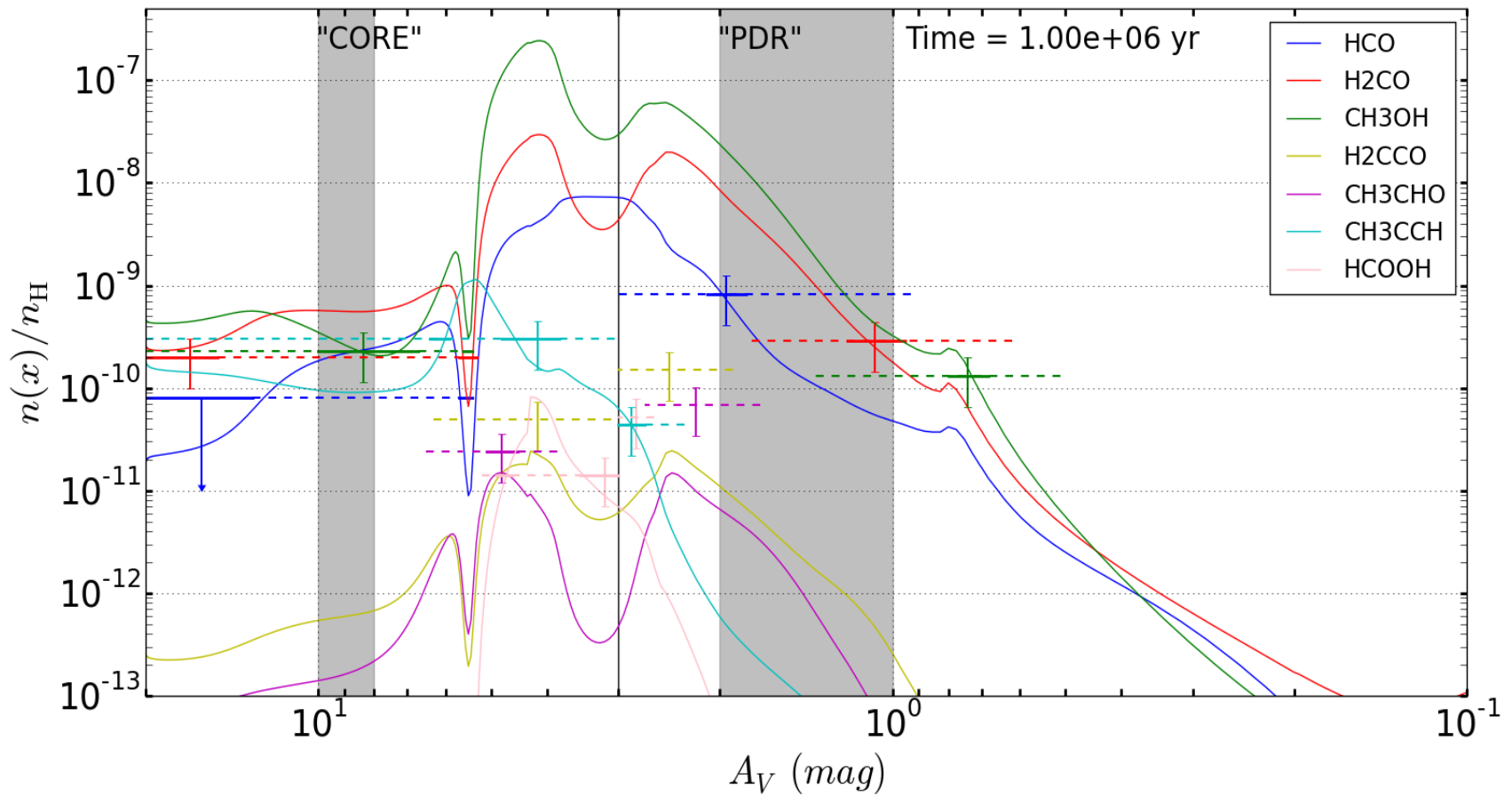
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# Physical structure

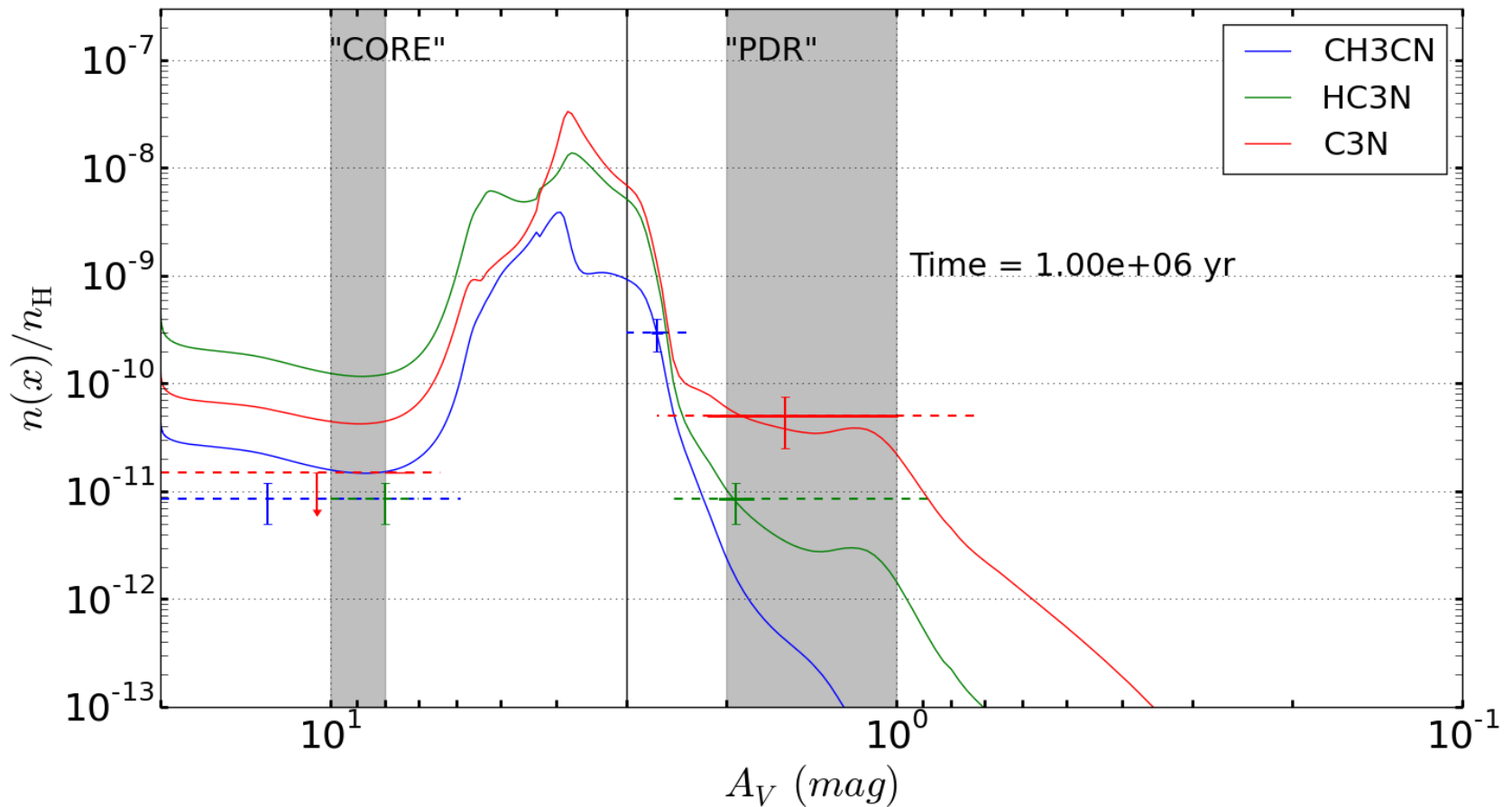


*Le Gal et al., A&A 605, A88 (2017)*

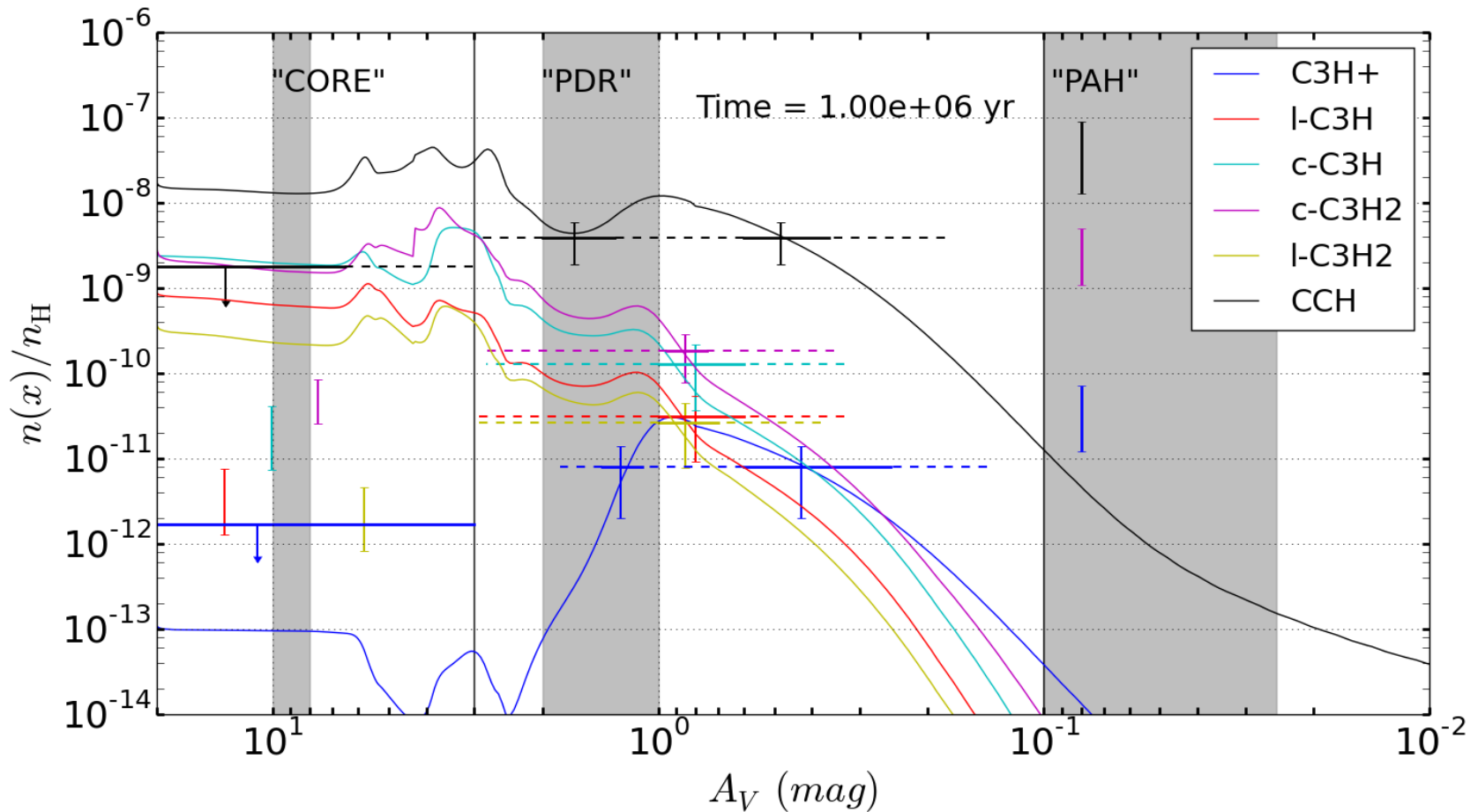
# Results: modeled abundances vs observations (I)



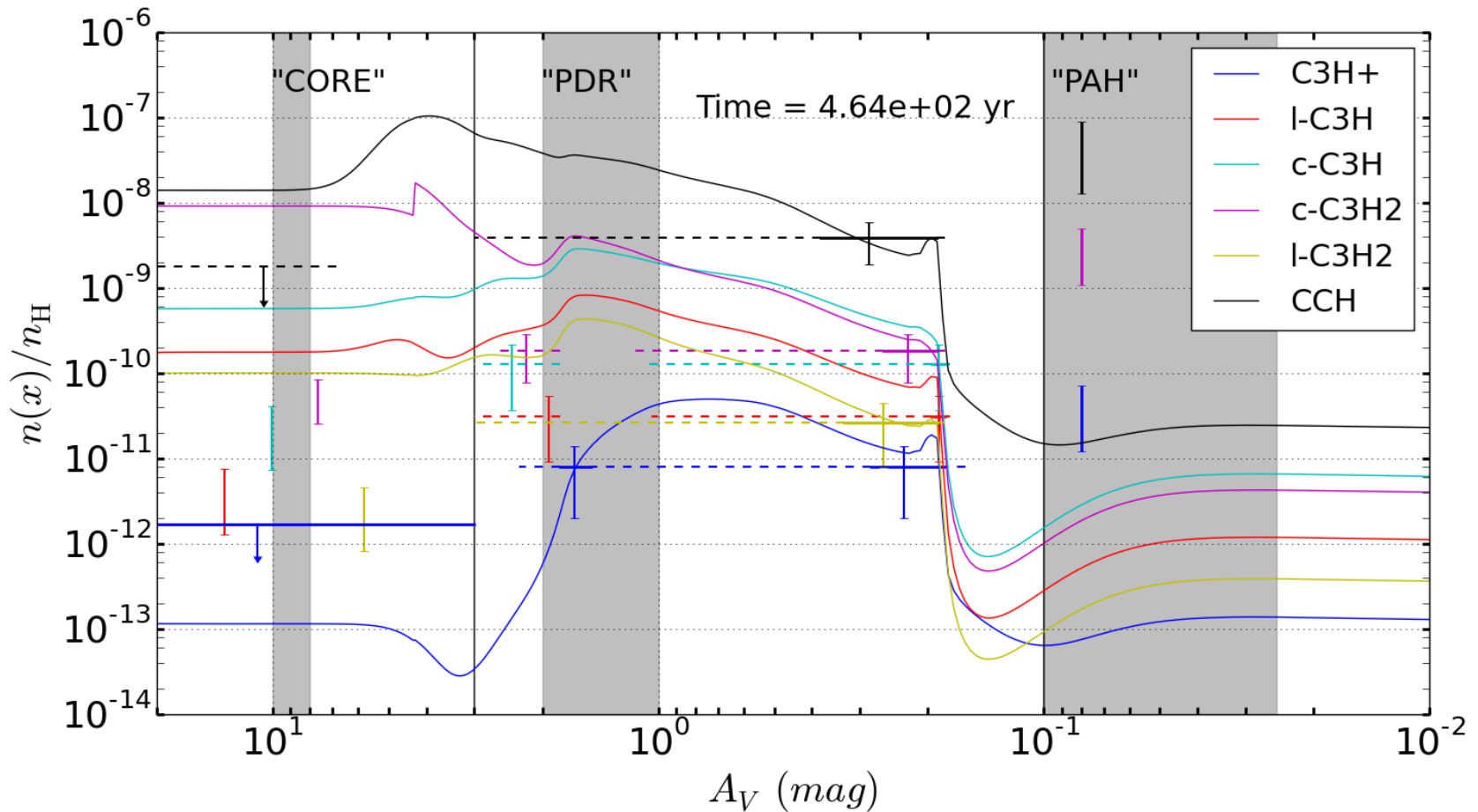
# Results: modeled abundances vs observations (II)



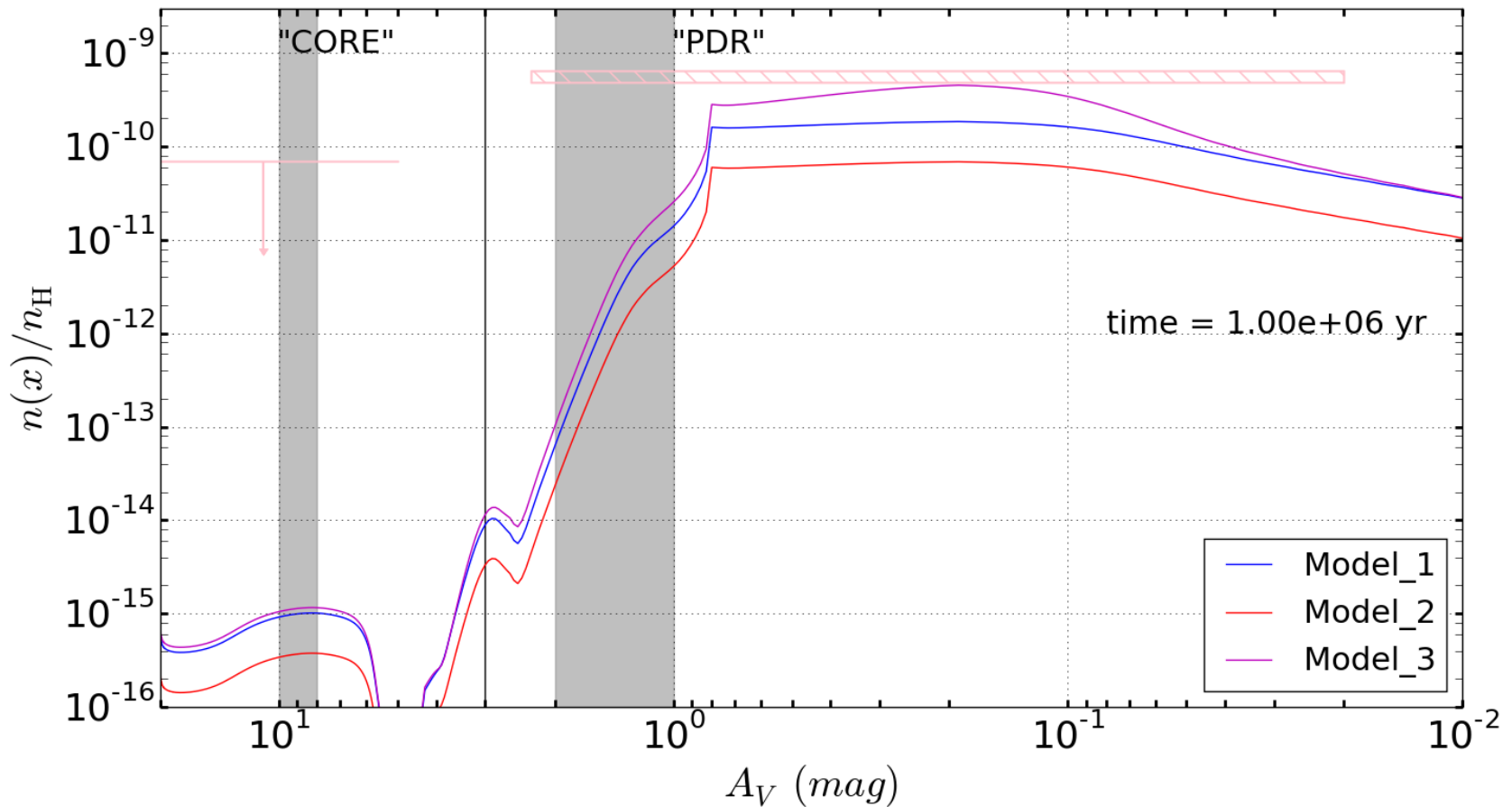
# Results: modeled abundances vs observations (III)



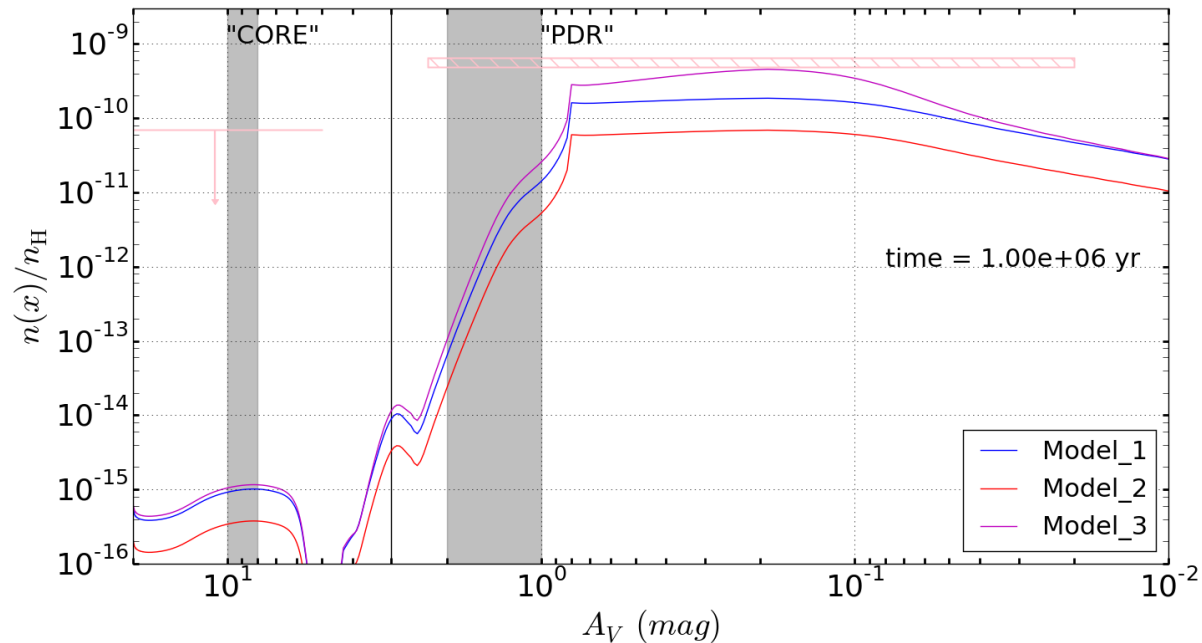
# Results: modeled abundances vs observations (IV)



# Results: modeled abundances vs observations (V)



# Results: modeled abundances vs observations (V)



- $\text{CF}^+ + \text{Photon} \longrightarrow \text{F} + \text{C}^+$  (Guzman et al. 2012),
- $\text{HF} + \text{C}^+ \longrightarrow \text{H} + \text{CF}^+$  (Neufeld et al. 2005; Guzman et al. 2012),
- $\text{CF}^+ + \text{e}^- \longrightarrow \text{C} + \text{F}$  (Novotny et al. 2005; Neufeld & Wolfire 2009; Guzman et al. 2012),
- $\text{F} + \text{H}_2 \longrightarrow \text{HF} + \text{H}$  (Tizniti et al. 2014).



# Summary & future works

## Summary:

### ➤ **Our time-dependent gas-grain chemistry model:**

- 0D model => initial starless molecular cloud = birth place of the  $\sigma$  Ori star, with Nautilus
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Thanks for  
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